Observational Signatures of First Stars

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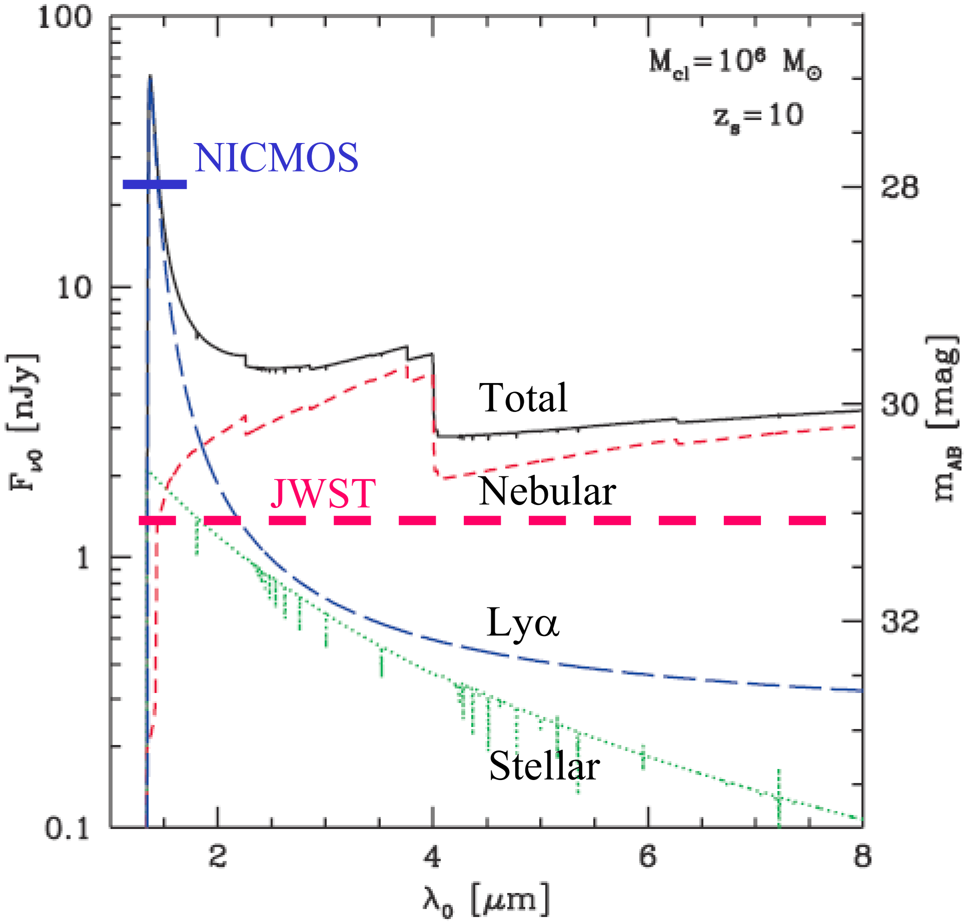
DIRECT DETECTABILITY

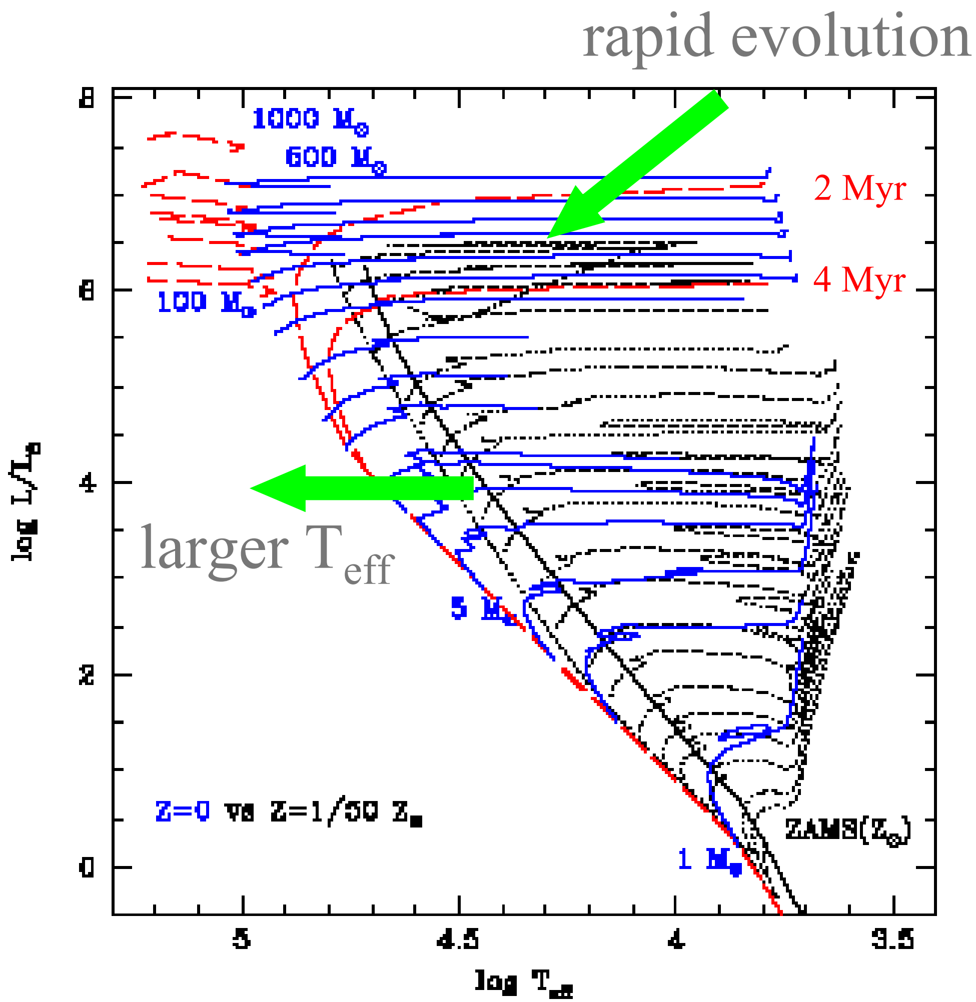


Pop III cluster

$M = 10^6\,M_\odot$
$z = 10$
$M_\star = 300\,M_\odot$

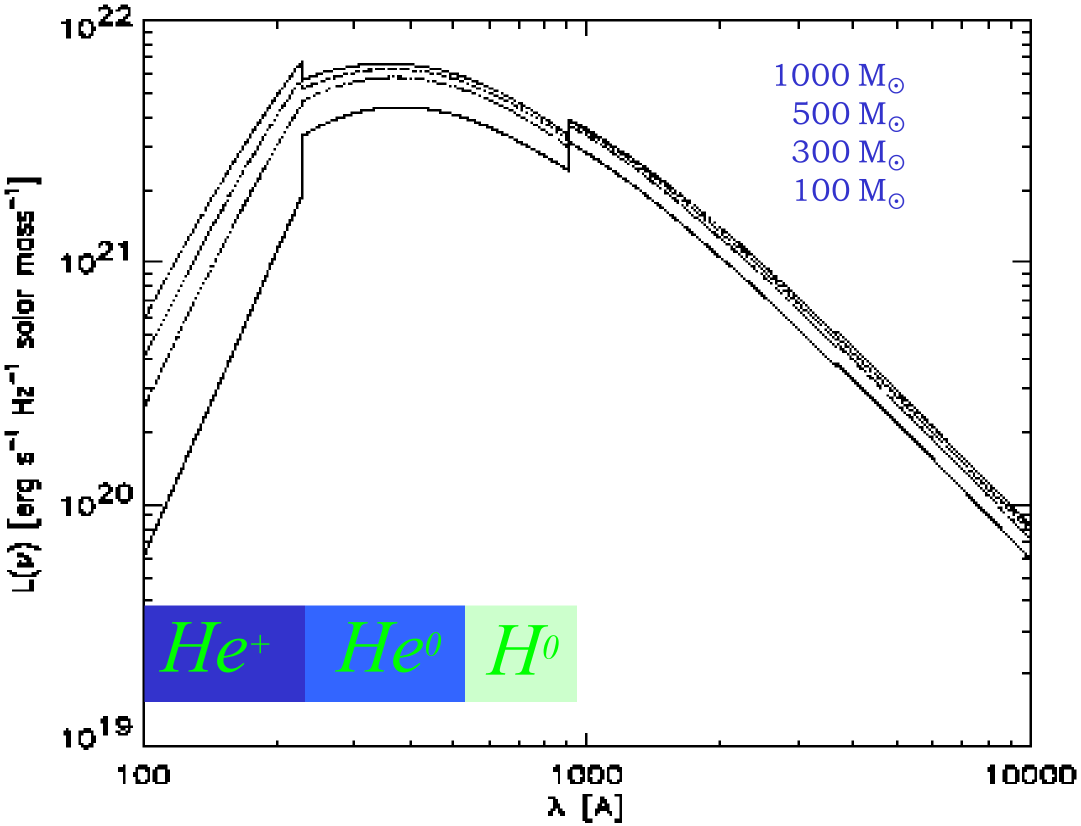
Schaerer 2002

STELLAR TRACKS



Bromm+ 2001

EMISSION SPECTRUM



IONIZING POWER

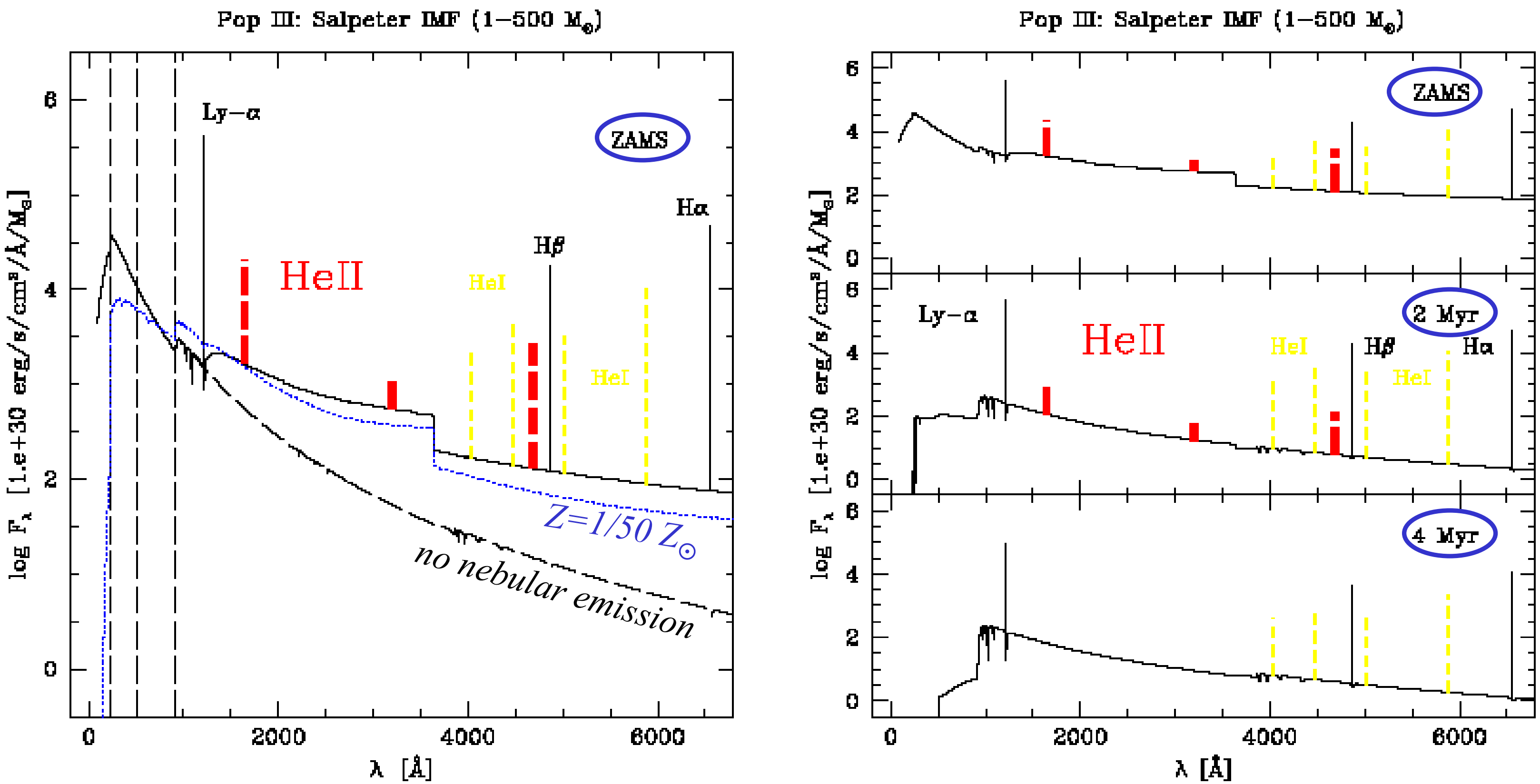
$$Q_i = 4\pi R_\star^2 q_i = 4\pi R_\star^2 \int_{\nu_i}^{\infty} \frac{F_\nu}{h\nu} d\nu, \qquad \bar{Q}_i(M) = \frac{\int_0^{t_\star(M)} Q_i(t, M) dt}{t_\star(M)},$$

Time-averaged quantities

$M_{\rm ini}$	lifetime	$\bar{Q}(\rm H)$	$\bar{Q}(\rm He^0)$	$\bar{Q}(\rm He^+)$	$\bar{Q}(\rm H_2)$	$\bar{Q}(\rm He^0)/\bar{Q}(\rm H)$	$\bar{Q}(\rm He^+)/\bar{Q}(\rm H)$
1000.		not available					
500.00	1.899E+06	6.802E+50	3.858E+50	5.793E+49	7.811E+50	0.567E+00	0.852E-01
400.00	1.974E+06	5.247E+50	3.260E+50	5.567E+49	5.865E+50	0.621E+00	0.106E+00
300.00	2.047E+06	3.754E+50	2.372E+50	4.190E+49	4.182E+50	0.632E+00	0.112E+00
200.00	2.204E+06	2.624E+50	1.628E+50	1.487E+49	2.918E+50	0.621E+00	0.567E-01
120.00	2.521E+06	1.391E+50	7.772E+49	5.009E+48	1.608E+50	0.559E+00	0.360E-01
80.00	3.012E+06	7.730E+49	4.317E+49	1.741E+48	8.889E+49	0.558E+00	0.225E-01
60.00	3.464E+06	4.795E+49	2.617E+49	5.136E+47	5.570E+49	0.546E+00	0.107E-01
40.00	3.864E+06	2.469E+49	1.316E+49	8.798E+46	2.903E+49	0.533E+00	0.356E-02
25.00	6.459E+06	7.583E+48	3.779E+48	3.643E+44	9.387E+48	0.498E+00	0.480E-04
15.00	1.040E+07	1.861E+48	8.289E+47	1.527E+43	2.526E+48	0.445E+00	0.820E-05
9.00	2.022E+07	2.807E+47	7.662E+46	3.550E+41	5.576E+47	0.273E+00	0.126E-05
5.00	6.190E+07	1.848E+45	1.461E+42	1.270E+37	6.281E+46	0.791E-03	0.687E-08
$R_{120/15}$	0.25	74	94	3×10^5	64	1.25	4390

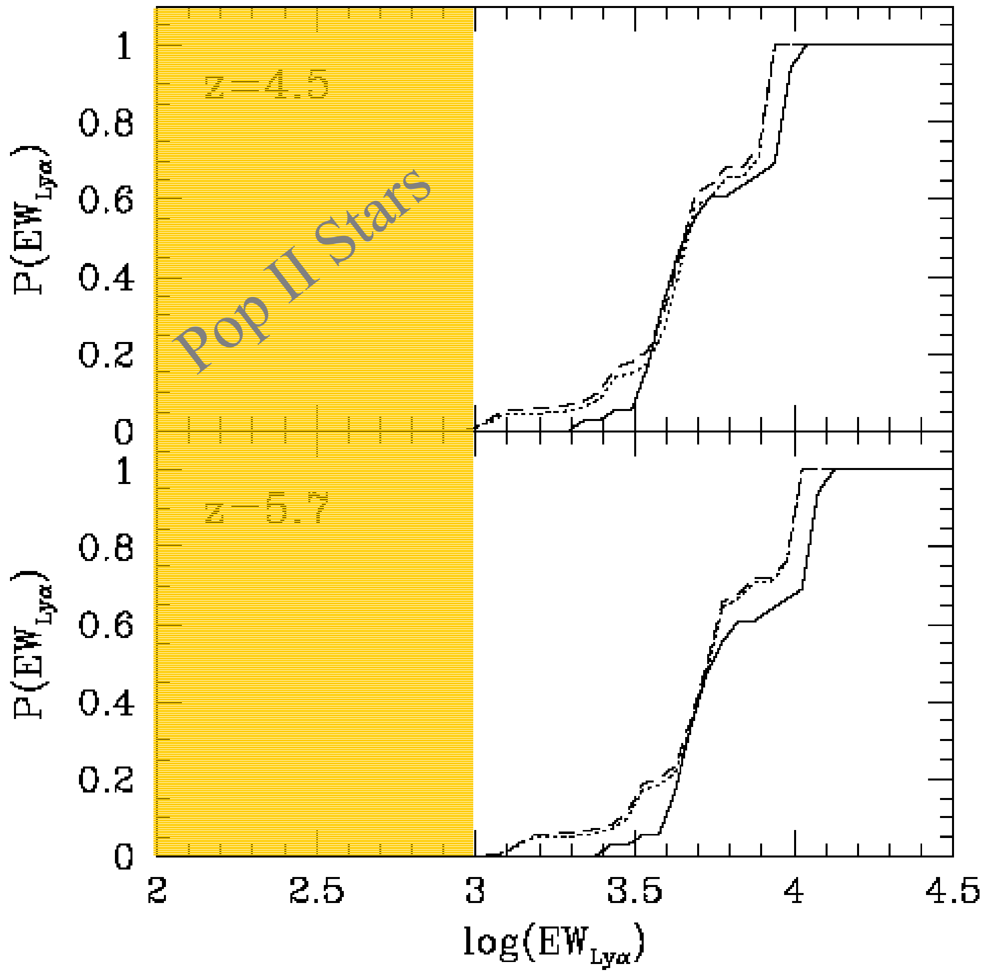
Schaerer 2002

HE NEBULAR LINES



Scannapieco+ 2002

LYA EW OF POP III STARS



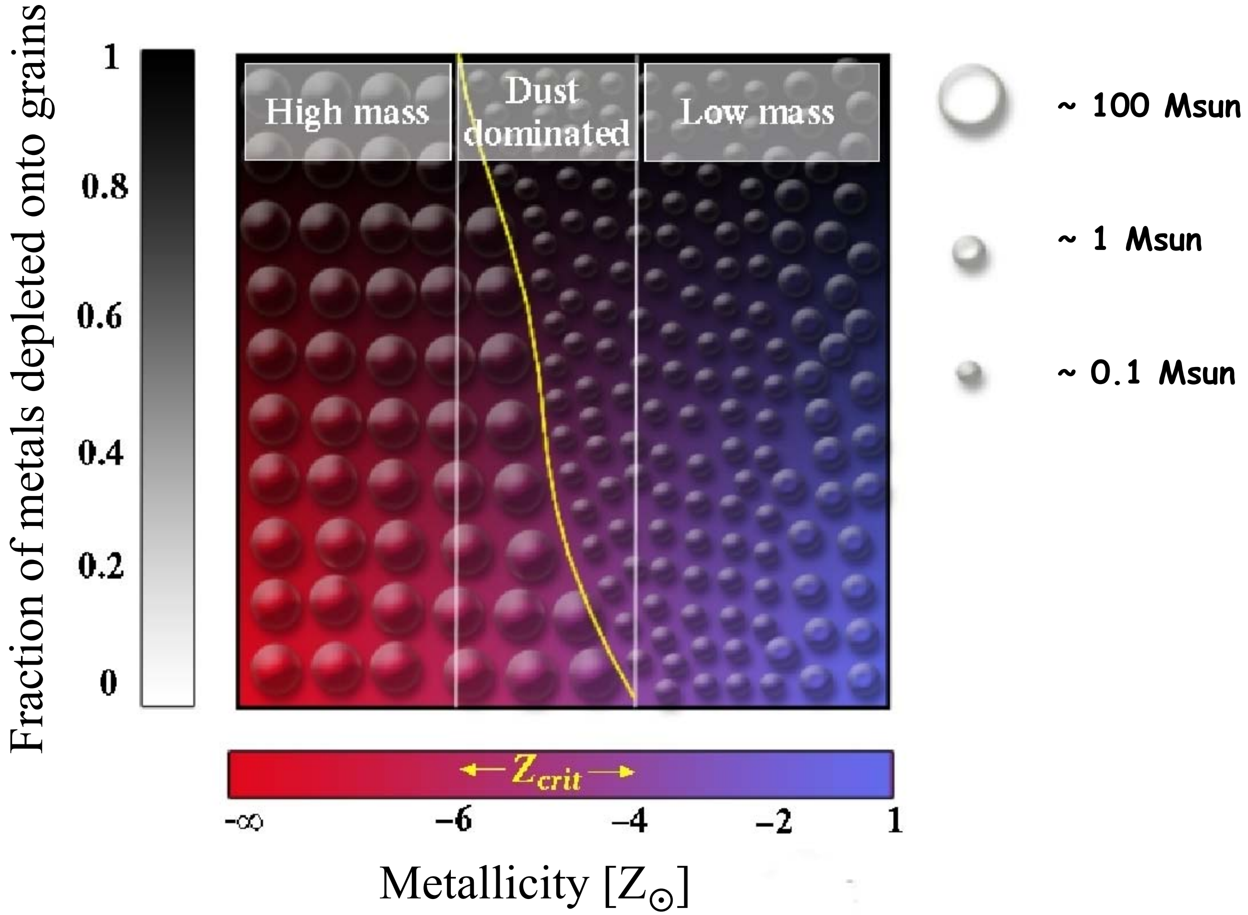
Pop III IMF

Salpeter 50-500 $M_\odot$

Salpeter 1-500 $M_\odot$

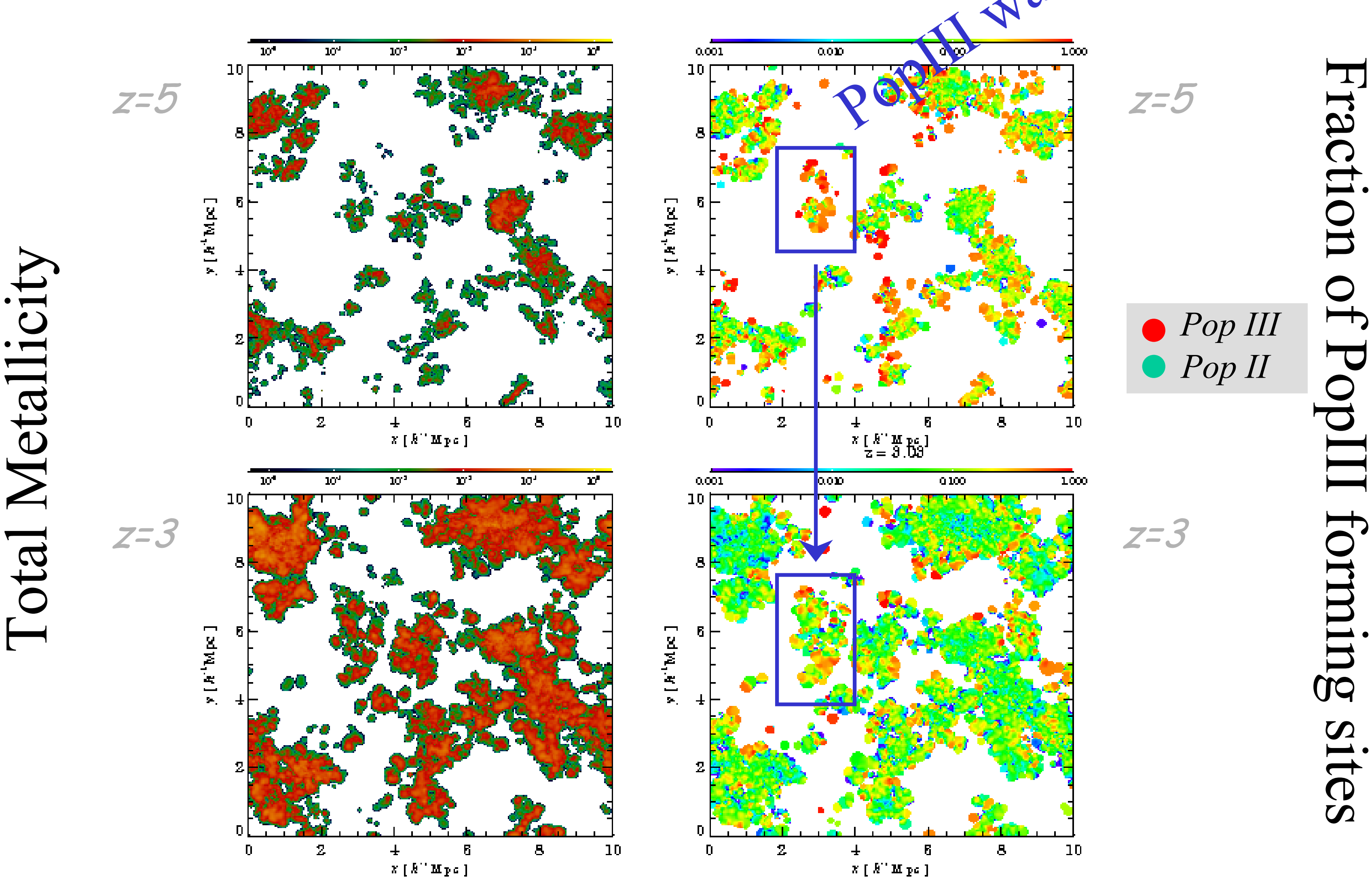
Schneider+ 2002, Schneider+ 2006

MASS OF EARLY STARS



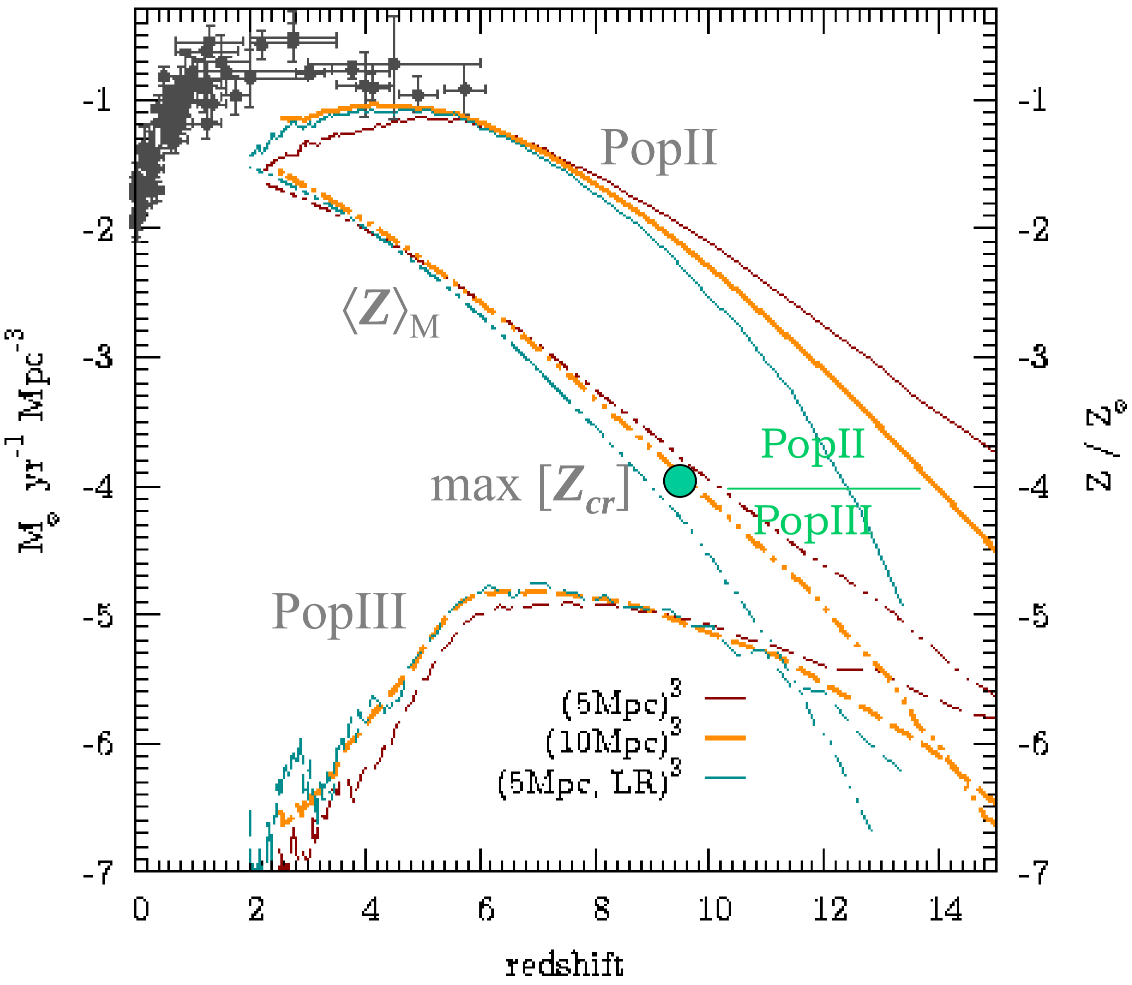
Tornatore, AF & Schneider 2007

COSMIC POPIII/POP II TRANSITION



Tornatore, AF & Schneider 2007

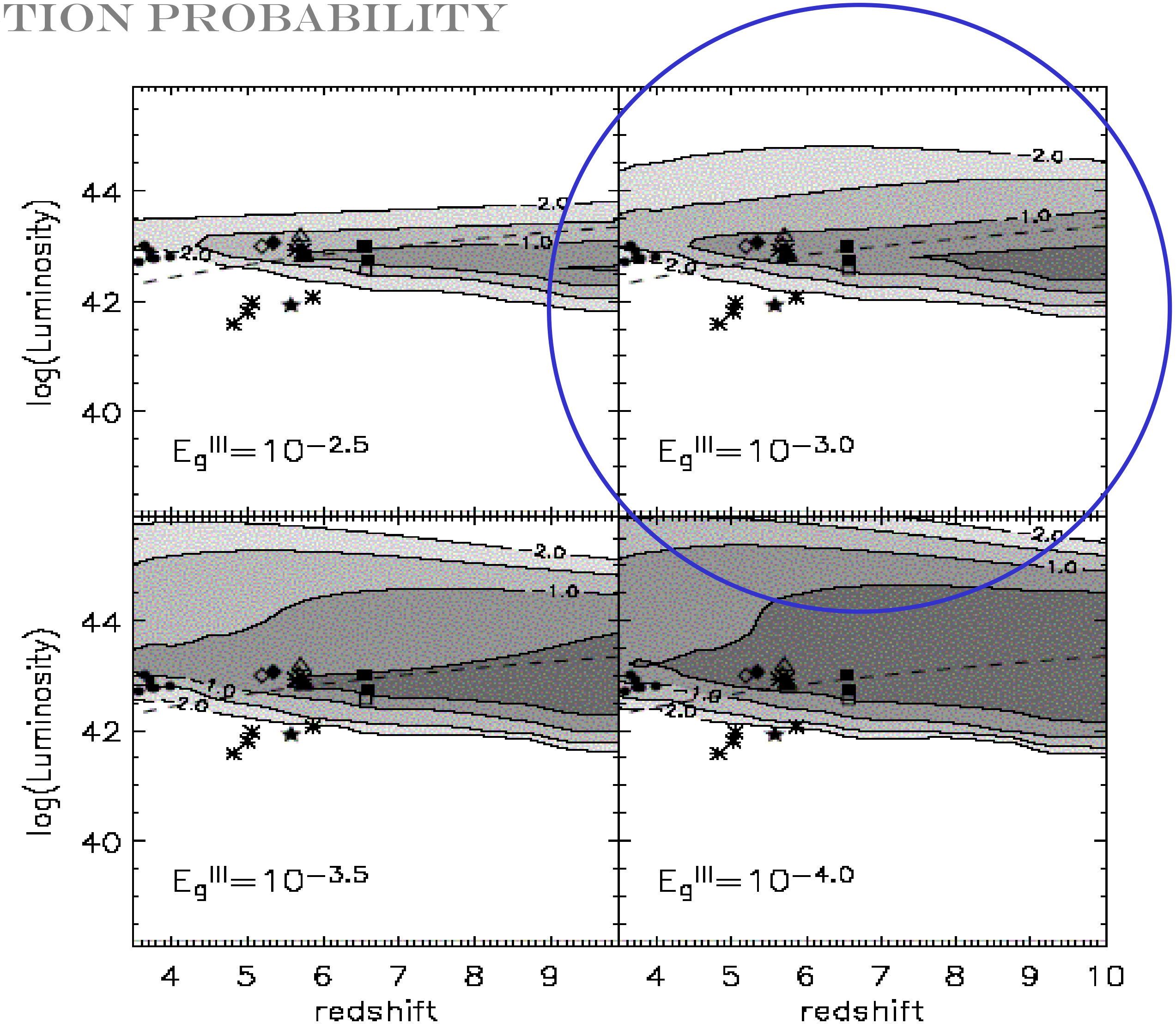
STAR FORMATION RATES



Scannapieco+ 2002

DETECTION PROBABILITY

Lyman Alpha Emitters



AVAILABLE SURVEYS

1. HK

Started 1980, 300 plates covering 2800 deg^2 (4100 deg^2) in the Northern (Southern) hemisphere, Plates visual inspection with binocular 10X microscope.Med-red (1 Å) spectroscopic follow-up with 2m-class telescopes

2. Hamburg/ESO Survey (HES)

Greatly increased numbers. Objective-prism survey, **2 mag deeper than HK***, regions of sky not sampled by HK. 8225 deg^2 above |b |=30^o. Selection using automatic spectral classification. Medium-res follow-up using 4m-class telescopes, JHK from 2MASS*

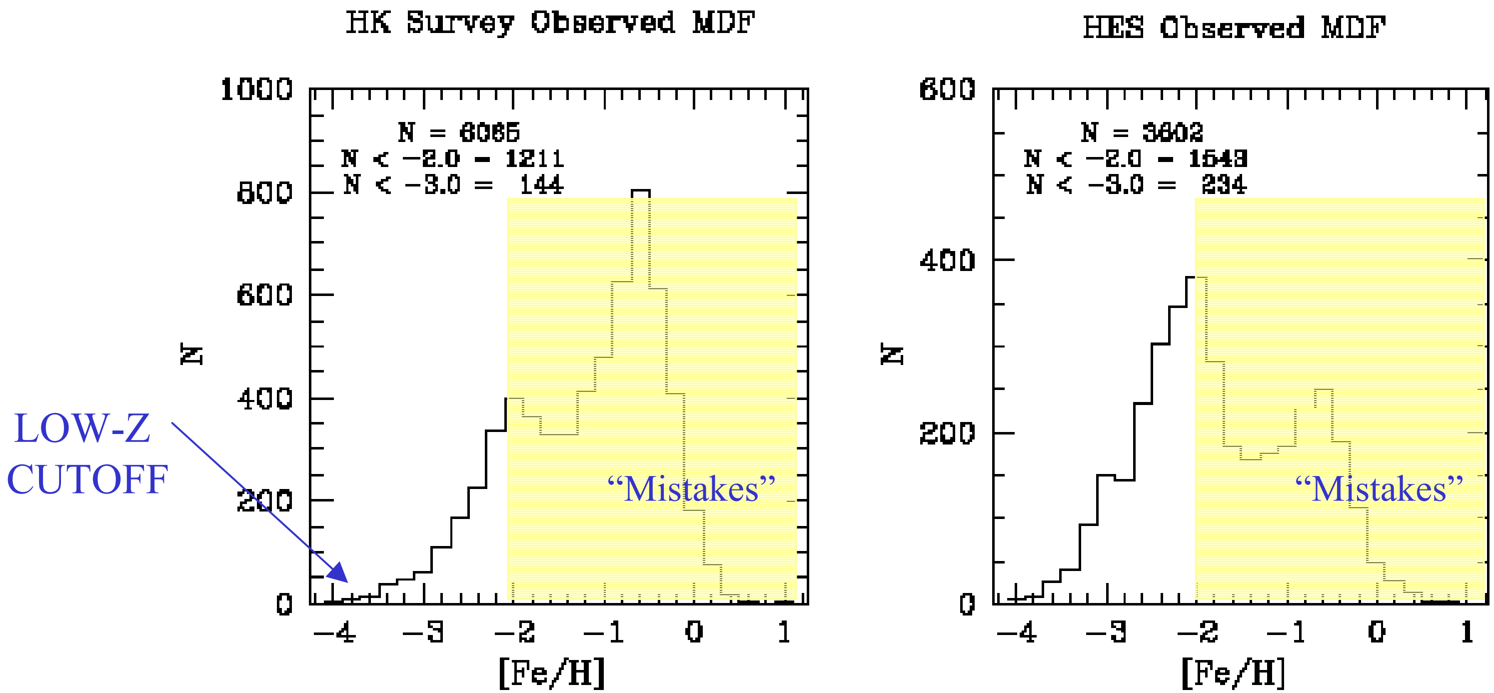
3. Sloan Digital Sky Survey (SDSS)

Med-res spectra + ugriz photometry for about 70000 stars, not targeted specifically to search for metal-poor stars, inhomogeneous assembly. Useful to test the tail of MDF

Table 1. Observational Follow-Up of Surveys

Survey	Spectra	Unique	UBV	JHK
HK	14488	11212	4944	10438
HES	7465	6212	812	5078
SDSS-DR3	71396	$\sim$ 70000	...	...

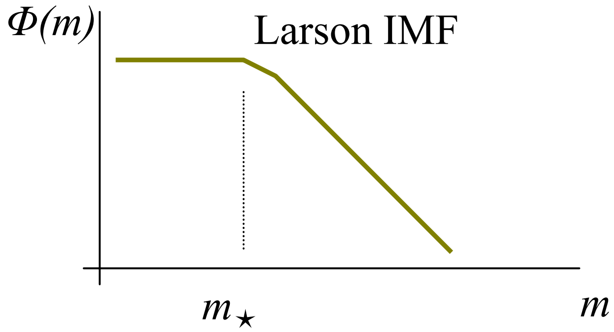
METALLICITY DISTRIBUTION FUNCTION

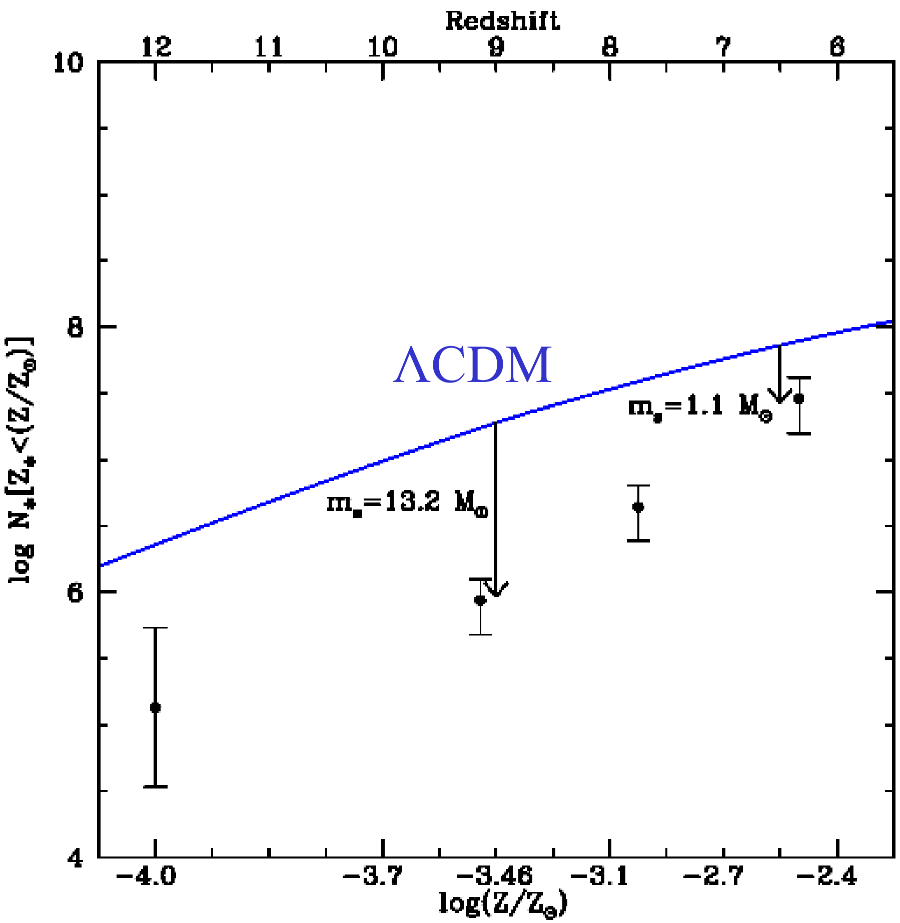


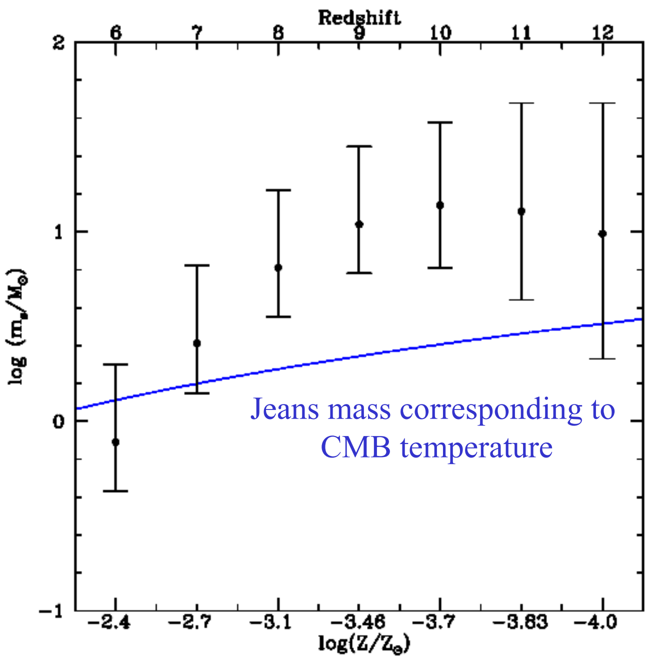
- *Total sample: 2700 stars with [Fe/H]<-2.0 and 400 with [Fe/H]<-3.0*
- *Bias due to selection criterion in the range –2.5 < [Fe/H]<-2.0*
- *Possible underestimate of the metallicity of cooler stars; spectrum-by-spectrum analysis.*

Hernandez & AF 2000, Tumlinson 2007

MDF INTERPRETED – I.



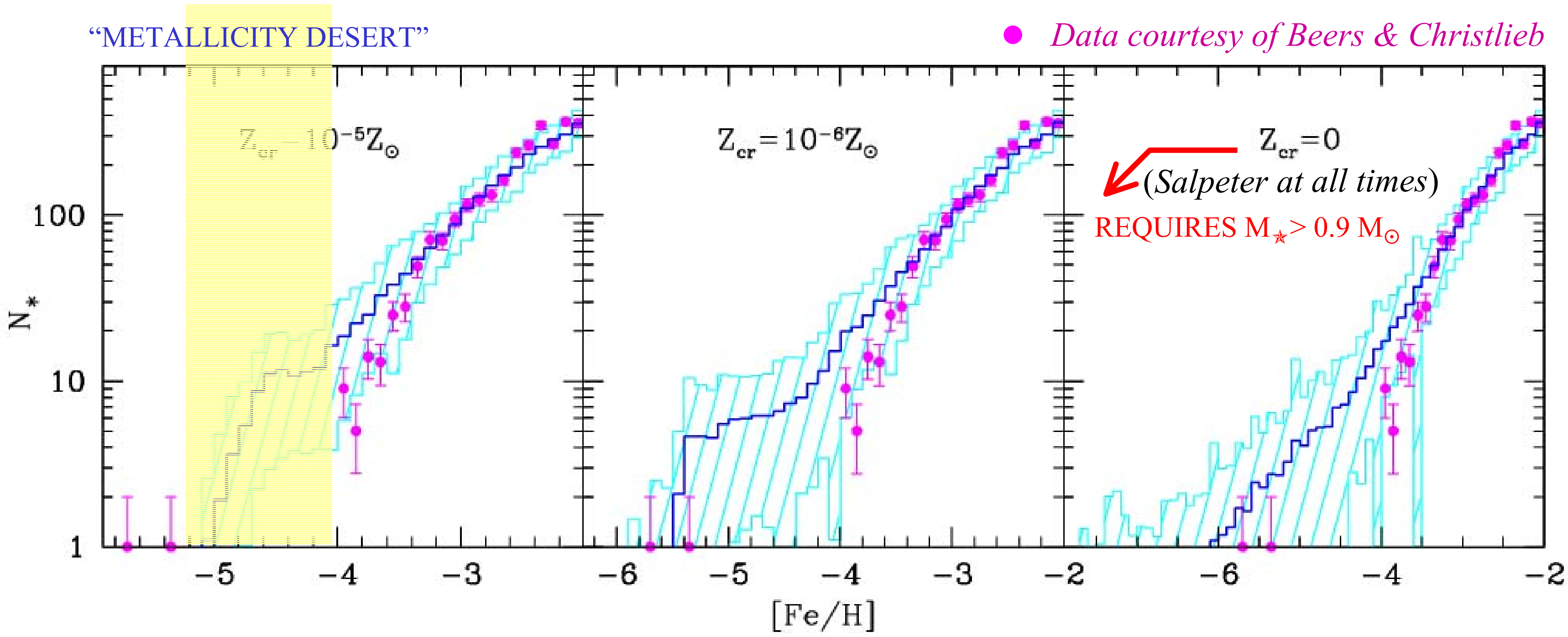




Salvadori, Schneider, AF 2006, Tumlinson 2006

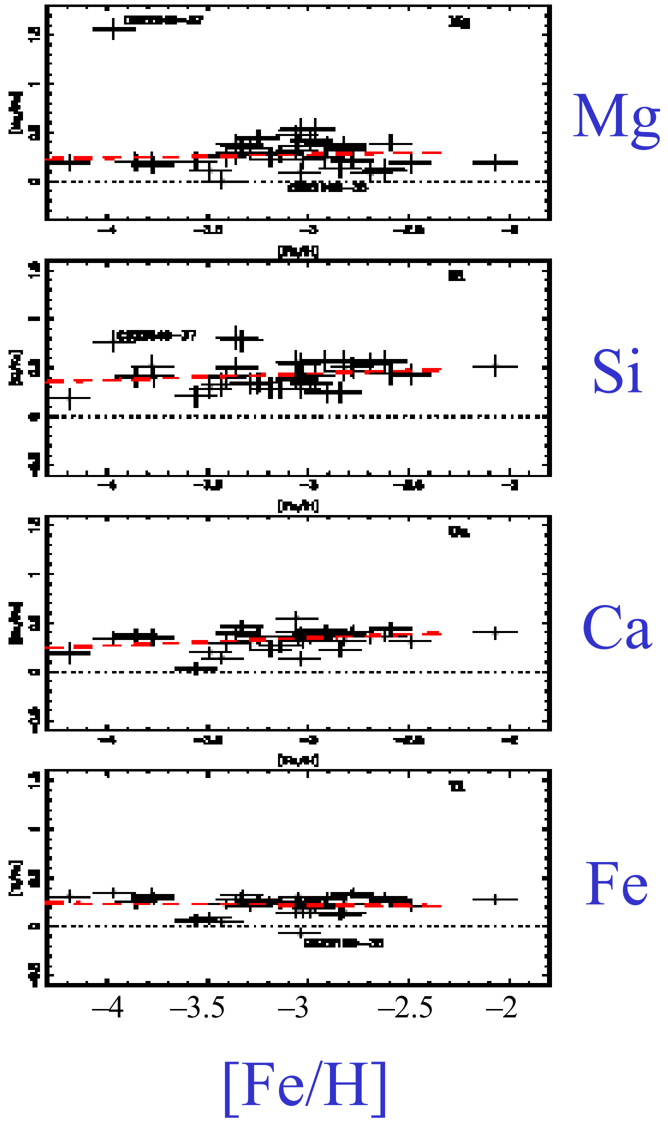
MDF INTERPRETED – II.

- ✓ Stellar / chemical evolution of the Milky Way based on ΛCDM merger-tree
- ✓ Joint HK/HES Metallicity Distribution Function, 2756 stars with [Fe/H] < – 2.



Cayrel+ 2004

ABUNDANCE PATTERNS



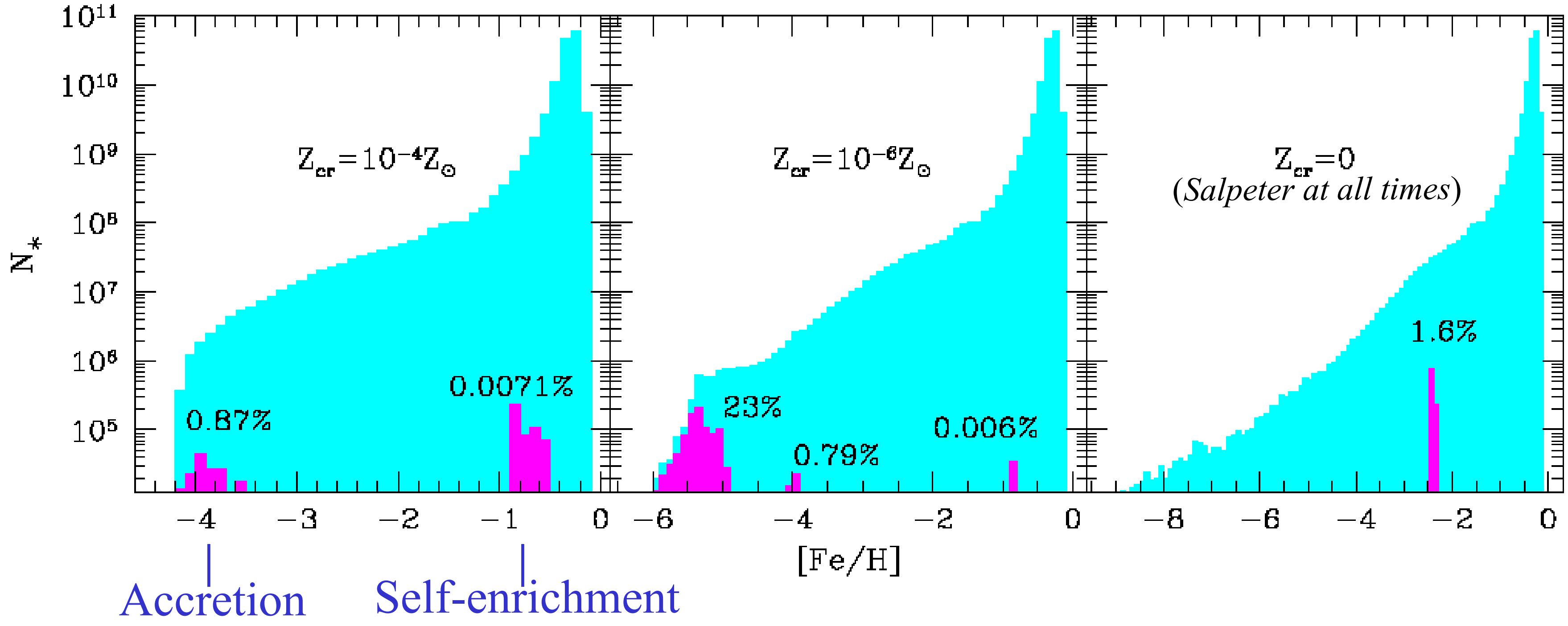
- ✓ 30/35 stars with –4.1 < [Fe/H] < –2.7
- ✓ 17 elements from C to Zn measured

Very small scatter: σ < 0.05 dex

- ▷ Unlikely resulting from *individual* SN ejecta
- ▷ Ratios *do not* match PISN yields

Salvadori, Schneider, AF 2006

SECOND GENERATION STARS

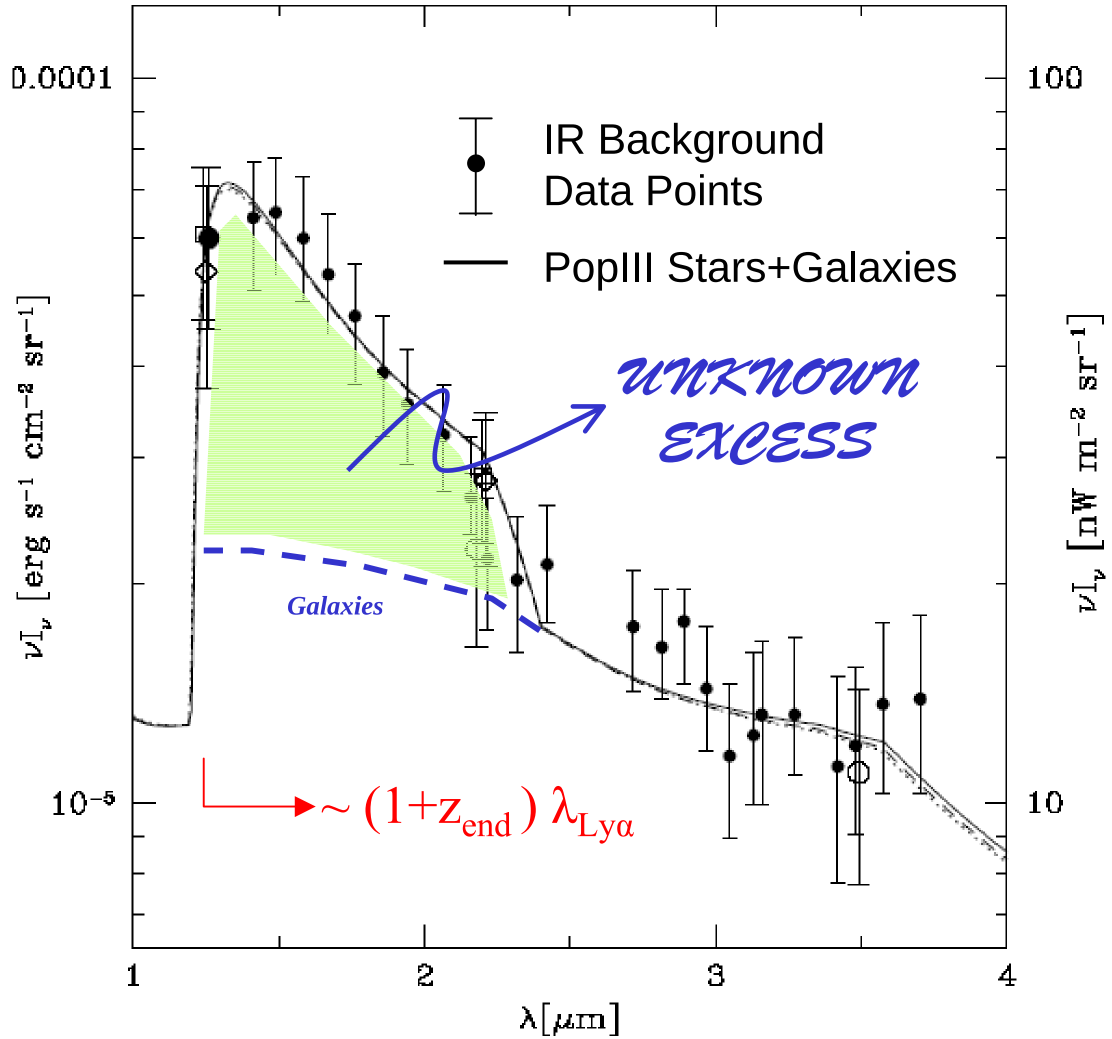


Joint HK/HES sample: 2756 stars with [Fe/H] < – 2.

Z_{crit}	Expected number of second-generation stars
$10^{-4}\ Z_{\odot}$	1.3
$10^{-6}\ Z_{\odot}$	0.3
0	0 06

Salvaterra & AF 2002, Santos+2002

A PUZZLING EXCESS



Best fit model to NIR data

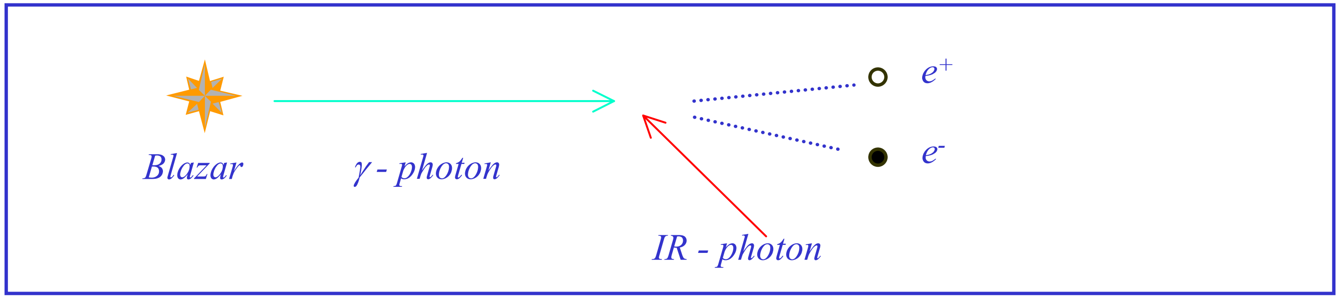
$z_{end} = 8.8$

$f_\star \approx 30\%$

Massive Pop III stars can explain NIRB excess

GAMMA-RAY CONSTRAINTS

- *TeV-GeV photons absorbed by optical/IR photons via $e^+ - e^-$ pair production.*



$$\tau(E) = \int_0^{z_{em}} dz \frac{dl}{dz} \int_{-1}^{1} dx \, \frac{(1-x)}{2} \int_{\epsilon_{th}}^{\infty} d\epsilon \, n(\epsilon) \, \sigma(\epsilon, E, x)$$

σ peaks at $\lambda_{IR} \sim 2.37$ (E/TeV) μm.

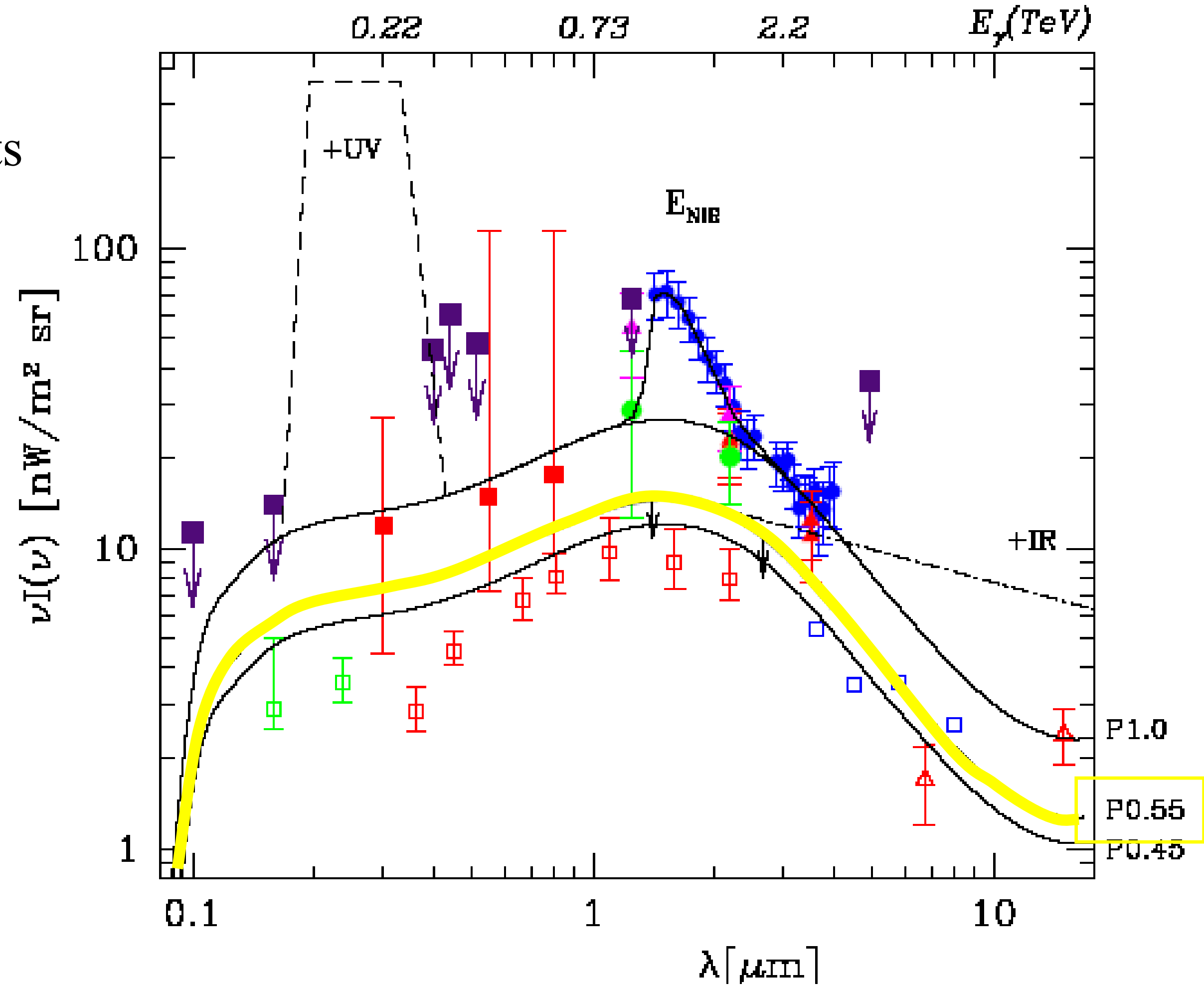
- *The observed spectrum of blazar reproduced by convolving the unabsorbed (power-law) spectrum with the optical depth:*

$$(dN/dE)_{abs} \propto e^{-\tau} E^{-\alpha}$$

Mapelli+ 2005, Aharonian+ 2005

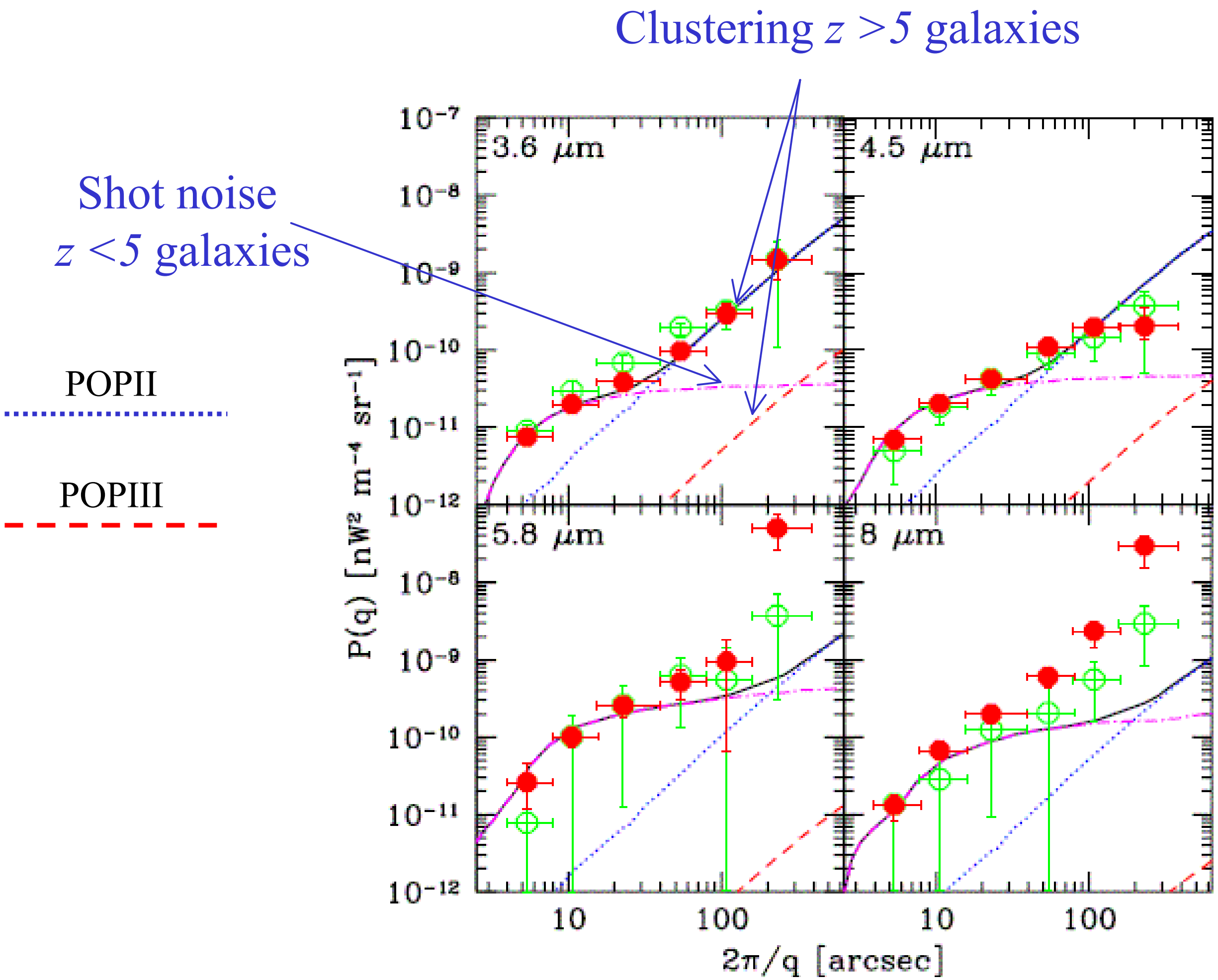
GAMMA-RAY CONSTRAINTS

- ○ Galaxy counts
- ● Direct measurements



Salvaterra+ 06, Cooray+ 06, Sullivan+06, Thompson+ 07a,b

FLUCTUATIONS

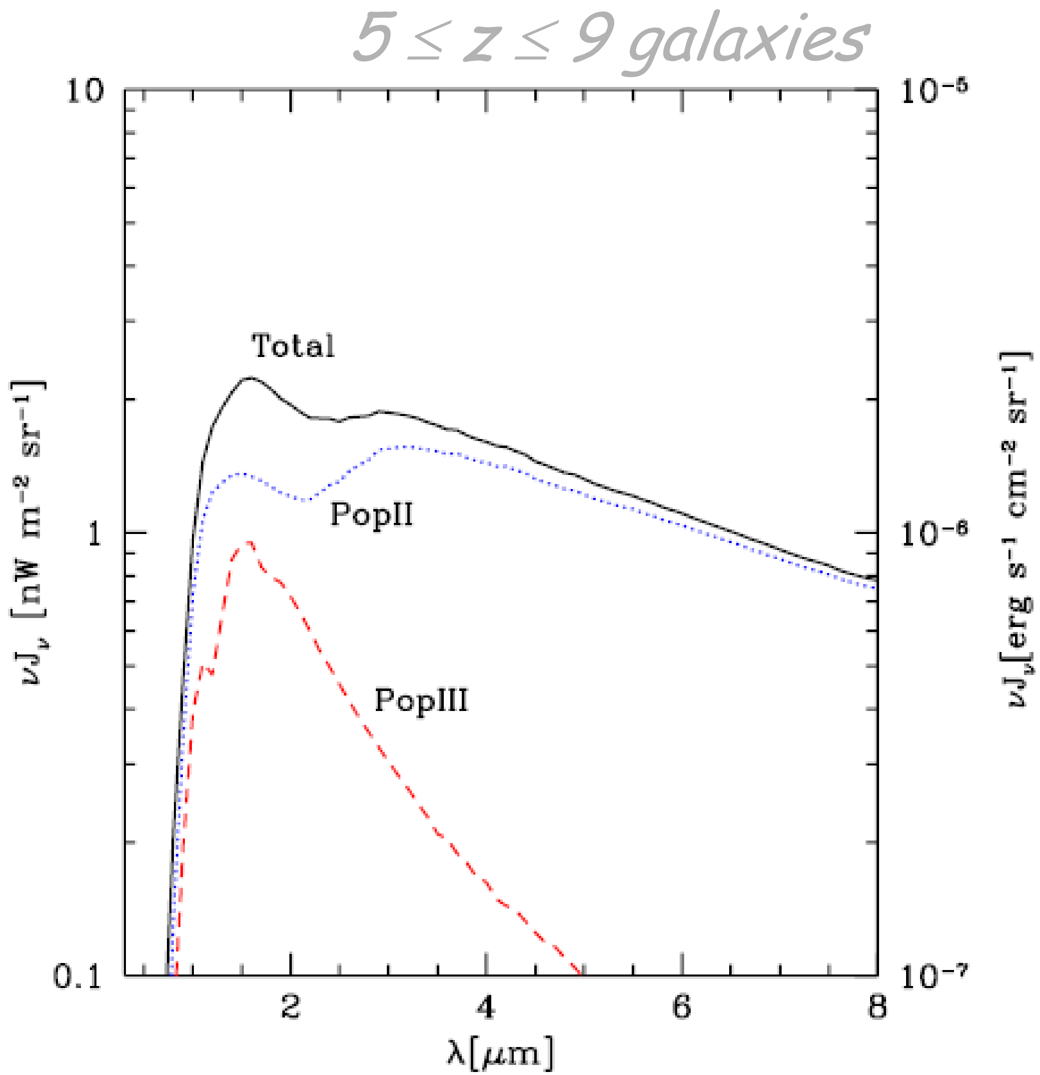


Spitzer/IRAC data

Kashlinsky+ 2005, 2007

Salvaterra+ 2006, Fernandez & Komatsu 2006

INTENSITY



NIRB PHOTON BUDGET

	nW m^{-2} sr^{-1} @ 1.4 μm
Observed	70
After zodi-subtraction (Wright)	17
Gamma rays	~15
Low-*z* galaxy contribution	> 8
Left unexplained	< 7
z>5 galaxies (from fluctuations)	2.5

SELECTED SUMMARY

- (Massive) PopIII stars strongly influence first stages of cosmic reionization
- Transition to normal stars occurs when $Z > Z_{crit} \sim 10^{-5\pm1} Z_{\odot}$; strongly governed by dust
- Pop III SF continues to $z \sim 3$ at periphery of collapsing structures. Observable in LAEs ?
- Metallicity Distribution Function of EMPs in the MW halo: hints on primordial IMF
- Imprint of very early (z>6) star formation activity left in the NIRB
- Experimental constraints on NIRB: intensity, fluctuations & pair-production opacity
- PISN explosions at moderate ($z < 6$) redshifts with JDEM

Scannapieco+ 2005, Weinmann & Lilly 2005

PISN EXPLOSIONS @ $z < 6$